

Fermion Stars with an Extra Dimension

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Abstract

Many efforts have been devoted to the studies of the phenomenology in particle physics with extra dimensions. We propose the degenerate fermion star in the five dimensions, and study what effects caused by the geometry of extra dimensions should appear in its structure. We note that Kaluza-Klein excited modes have effects for the larger scale of extra dimensions and examine the conditions on which different layers should be caused in the inside of the stars. We expound how the effects of the extra dimensions appears on physical quantities.

1 Introduction

Recently, physics dealing with higher dimensions has been noted [1, 2]. Considering extra dimensions in GUT, one points out four forces could be unified at low energy scale. And these are closely related with superstring theories, supergravity, and the others. There are many works on extra dimensions. One of them is “Neutron stars and extra dimensions” by Liddle, Moorhouse and Henriques [3]. They considered the neutron stars in five dimensions of which the fifth one is compactified into S^1 , and showed that the maximum mass of the stars decreases under the background. However there are some problems in their result. First, they did not consider Kaluza-Klein (K-K) excited modes caused by the geometry of extra dimensions. Secondly, they put some hypotheses both on the equation of state and on that of conservation, so arbitrariness is left. Whereas, in this paper, taking account of the K-K modes, we propose the star in the five dimensions, which is made of degenerate Dirac fermions, and indicated that the excited modes have effects to the inside of star. After asking for the maximum mass and the radius by use of numerical calculation, we will make it clear what the interior structure of star should be.

2 Matter

2.1 (3+1) dimensions

In the four dimensional theory, the thermodynamic potential of a fermion gas with mass m and half spin is

$$\Omega_4 = -2 \frac{1}{\beta} V \int \frac{d^3 \mathbf{p}}{(2\pi)^3} \left[\ln \left(1 + e^{-\beta(\sqrt{\mathbf{p}^2 + m^2} - \mu)} \right) + (\mu \leftrightarrow -\mu) \right], \quad (1)$$

where we put subscript to emphasize that it stands for a four dimensional quantity. We will restrict our system to degenerate fermion gas and take the zero temperature limit. Then, the thermodynamic potential becomes

$$\Omega_4(m) \equiv \begin{cases} -V \frac{m^4}{24\pi^2} \left[\frac{\mu}{m} \sqrt{\frac{\mu^2}{m^2} - 1} \left(2 \frac{\mu^2}{m^2} - 5 \right) + 3 \ln \left| \frac{\mu}{m} + \sqrt{\frac{\mu^2}{m^2} - 1} \right| \right] & (m < \mu) \\ 0, & (m \geq \mu) \end{cases} \quad (2)$$

Thermodynamical quantities follow from Eq.(2).

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2.2 (4+1) dimensions

We will extend our argument into the five dimensional theory. Here we suppose that the fifth dimension is compactified into S^1 with a radius b , b is not so large. We impose the periodic boundary condition on a wave function in the fifth dimension:

$$\psi(\chi) \sim e^{ip_5 \cdot \chi} , \quad (3)$$

$$\psi(\chi + 2\pi b) \sim \psi(\chi) , \quad (4)$$

with the fifth coordinate χ and momentum p_5 , then

$$p_5 = \frac{n}{b} \quad (n: \text{integer}) . \quad (5)$$

Therefore the relativistic energy involving the fifth dimension is

$$E_5 = \sqrt{p^2 + \left(\frac{n}{b}\right)^2 + m^2} \quad (6)$$

$$= \sqrt{p^2 + M^2} , \quad (7)$$

where

$$M^2 \equiv \left(\frac{n}{b}\right)^2 + m^2 , \quad (8)$$

(n is integer. Unless $n = 0$, the K-K modes become effective.) Using Eq.(7), the thermodynamic potential of the fermion gas with mass m and half spin in the (4+1) dimensional space-time is

$$\Omega_5 = -2\frac{1}{\beta}V_4 \int \frac{d^4 p}{(2\pi)^4} \left[\ln \left(1 + e^{-\beta(\sqrt{p^2 + M^2} - \mu)} \right) + (\mu \leftrightarrow -\mu) \right] , \quad (9)$$

If we turn the integral over the fifth dimensional momentum into the sum over quantum number n , that is :

$$\int dp_5 \rightarrow \frac{1}{b} \sum_n , \quad (10)$$

and in addition,

$$V_4 = 2\pi b V , \quad (11)$$

then the thermodynamic potential is

$$\begin{aligned} \Omega_5 &= -2\frac{1}{\beta}V \sum_n \int \frac{d^3 p}{(2\pi)^3} \left[\ln \left(1 + e^{-\beta(\sqrt{p^2 + \frac{n^2}{b^2} + m^2} - \mu)} \right) + (\mu \leftrightarrow -\mu) \right] \\ &= \sum_n \Omega_4 \left(\sqrt{m^2 + \frac{n^2}{b^2}} \right) = \sum_n \Omega_4(M) . \end{aligned} \quad (12)$$

As well as the previous section, we will deal with the degenerate fermion gas and low temperature near to zero. The thermodynamic potential, therefore,

$$\Omega_5 = \begin{cases} -Vb^{-4}f(\mu b, mb), & (\mu > M) \\ 0, & (\mu \leq M) \end{cases} , \quad (13)$$

where

$$f \equiv \sum_{n=-\infty}^{\infty} \frac{(Mb)^4}{24\pi^2} \left\{ \frac{\mu b}{Mb} \left(2 \left(\frac{\mu b}{Mb} \right)^2 - 5 \right) \sqrt{\left(\frac{\mu b}{Mb} \right)^2 - 1} + 3 \ln \left| \sqrt{\left(\frac{\mu b}{Mb} \right)^2 - 1} + \frac{\mu b}{Mb} \right| \right\} \quad (\mu > M). \quad (14)$$

From Eq.(13), thermodynamical quantities are as follows:

$$P = -\frac{1}{2\pi b} \frac{\partial \Omega_5}{\partial V} = -\frac{1}{2\pi b} \frac{\Omega_5}{V} = \frac{1}{2\pi b} \frac{1}{b^4} f, \quad (15)$$

$$P_5 = -\frac{1}{2\pi V} \frac{\partial \Omega_5}{\partial b} = \frac{1}{2\pi b} \frac{1}{b^4} \left(x \frac{\partial f}{\partial x} + y \frac{\partial f}{\partial y} - 4f \right), \quad (16)$$

$$\rho = \frac{U}{V_4} = \frac{1}{2\pi b} \frac{1}{b^4} \left(x \frac{\partial f}{\partial x} - f \right), \quad (17)$$

where P_5 is the pressure in the fifth dimension, and we put $x = \mu b$, $y = mb$.

3 Space-Time

We take the line element to be of the form

$$ds^2 = e^{-\Phi} \left[-e^{-2\delta} \Delta dt^2 + \frac{dr^2}{\Delta} + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2) \right] + b_0^2 e^{2\Phi} d\chi^2, \quad (18)$$

where we regard Δ , δ , and Φ ($b = b_0 e^{\Phi}$) as functions depending only on r , the distance from the origin. Next, the energy-momentum tensor is

$$T_{\nu}^{\mu} = \text{diag.}(-\rho, P, P, P, P_5), \quad (19)$$

in which we suppose isotropic pressure in the space of three dimensions, and represent the fifth dimensional pressure as P_5 . The equation of conservation is

$$\nabla_{\mu} T^{\mu r} = 0. \quad (20)$$

From this, we find the chemical potential μ satisfies that

$$\mu b = \frac{e^{\frac{1}{2}\Phi + \delta}}{\sqrt{\Delta}} \mu_0 b_0, \quad (21)$$

where μ depends on r and μ_0 and b_0 are each the value of μ and b when r is close to zero.

4 Equations

Now, we are just deriving Einstein equations. Before that, we will rewrite Eq.(14). After putting

$$x = \mu b, \quad y = mb,$$

Eq.(14) reduces to

$$\begin{aligned} f(x, y) &= \sum_{n=-\infty}^{\infty} \frac{(y^2 + n^2)^2}{24\pi^2} \left[\frac{x}{\sqrt{y^2 + n^2}} \sqrt{\frac{x^2}{y^2 + n^2} - 1} \left(2 \frac{x^2}{y^2 + n^2} - 5 \right) \right. \\ &\quad \left. + 3 \ln \left| \frac{x}{\sqrt{y^2 + n^2}} + \sqrt{\frac{x^2}{y^2 + n^2} - 1} \right| \right] \\ &= \sum_{n=-\infty}^{\infty} \tilde{f}(x, y) \quad (\sqrt{x^2 - y^2} > n^2), \end{aligned} \quad (22)$$

where the sum over n is done unless n^2 exceeds $\sqrt{x^2 - y^2}$. (To remark this, we put prime on a sum symbol.) In addition, we put

$$Y = \sqrt{y^2 + n^2}. \quad (23)$$

Using the leading formulae, Einstein equations are

$$\frac{\tilde{M}'_\star}{\tilde{r}^2} - \frac{3}{8}\tilde{\Delta}(\Phi')^2 = 4\pi \frac{1}{m^4 b_0^4} e^{-6\Phi} \sum' \left(x \frac{\partial \tilde{f}}{\partial x} - \tilde{f} \right), \quad (24)$$

$$\frac{1}{\tilde{r}}\delta' + \frac{3}{4}(\Phi')^2 = -4\pi \frac{1}{m^4 b_0^4} \frac{e^{-6\Phi}}{\tilde{\Delta}} \sum' \left(x \frac{\partial \tilde{f}}{\partial x} \right), \quad (25)$$

$$\begin{aligned} \Phi'' + \left(\frac{\tilde{\Delta}'}{\tilde{\Delta}} - \delta' + \frac{2}{\tilde{r}} \right) \Phi' &= -\frac{8\pi}{3} \frac{1}{m^4 b_0^4} \frac{e^{-6\Phi}}{\tilde{\Delta}} \sum' \left(\frac{3Y^2 - 2y^2}{Y^2} x \frac{\partial \tilde{f}}{\partial x} \right. \\ &\quad \left. - \frac{12Y^2 - 8y^2}{Y^2} \tilde{f} \right), \end{aligned} \quad (26)$$

where

$$\tilde{M}_\star = \sqrt{G_0^{(4)} m^2 G_0^{(4)}} m M_\star, \quad (27)$$

$$\tilde{r} = \sqrt{G_0^{(4)} m^4} r, \quad (28)$$

$$\tilde{\Delta} = 1 - \frac{2\tilde{M}_\star}{\tilde{r}}. \quad (29)$$

In the above equations, we define a usual Newtonian constant $G_0^{(4)}$ as

$$G_0^{(4)} \equiv \frac{G_0^{(5)}}{2\pi b_0}, \quad (30)$$

in which $G_0^{(5)}$ stands for the Newtonian constant in the fifth dimension. And M_\star is the mass of star depending on r . A prime means derivative with respect to \tilde{r} . We can solve these equations numerically. In the next section, we will show the result.

5 Results

We will exhibit the relationship between the mass and the central density in fermion stars in Fig.1, and the mass and the radius in Fig.2. It turns out that the larger the extra dimension grows, the smaller the maximum mass becomes. Furthermore, it is remarkable that two maximum points appeared when $mb_0 = 3, 4$.

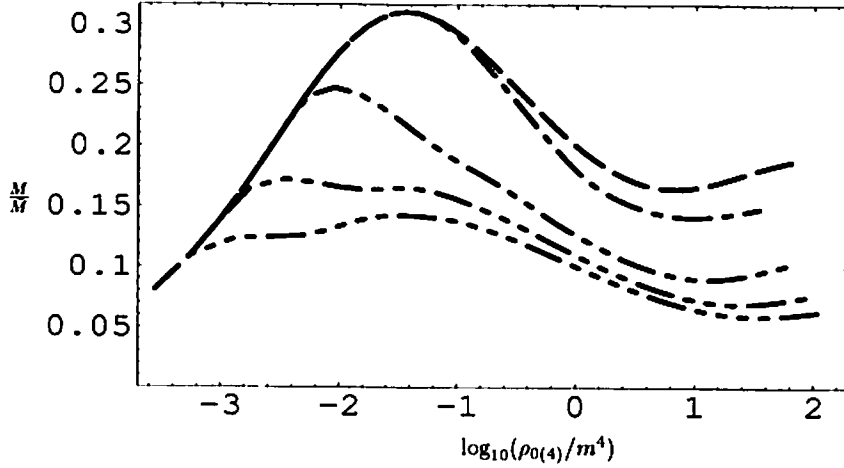


Fig 1: Plots of the mass M of the fermion stars versus its central density $\rho_{0(4)}$ for the various scales of the extra dimension. The dashed line corresponds to $mb_0 \approx 0$. The dot-dashed line corresponds to $mb_0 = 1$, the two dot-dashed line to $mb_0 = 2$, the three to $mb_0 = 3$, the four to $mb_0 = 4$.

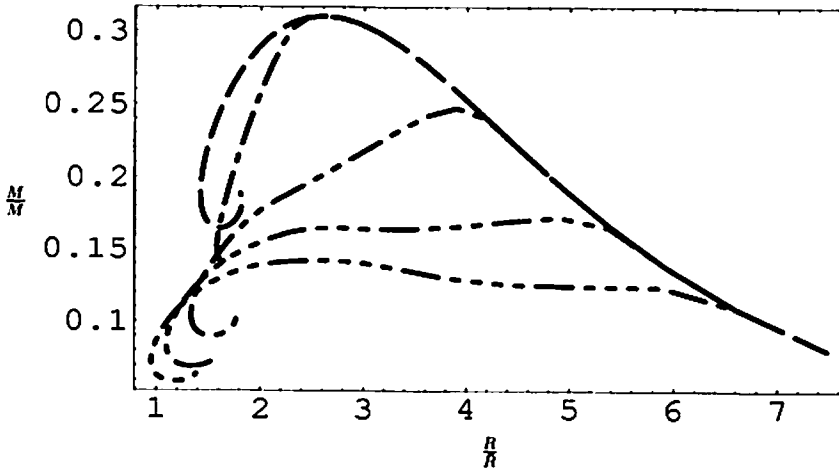


Fig 2: Plots of the mass M of the fermion stars versus its radius R for the various scales of the extra dimension. The correspondence between the dashed lines and the fifth dimension is the same as Fig.1.

Next, in Fig.3, we draw the interior structure of the stars having the maximum mass, which is based on Fig.2. We can read off it from Fig.3 that the excited modes have effects in the core of the stars with increase in the extra dimension. We have two solutions in the case of $mb_0 = 3, 4$. For $mb_0 = 3$, one of them is a larger star than that for $mb_0 = 2$ and its central density is lower, which is shown in Fig.1. While the other gets to be a smaller one and its central density is higher. In this solution, higher excited mode ($n=2$) is caused mainly in the center of stars and lower mode ($n=1$) in the vast region including the core. Proceeding to $mb_0 = 4$, this inclination appears more remarkably.

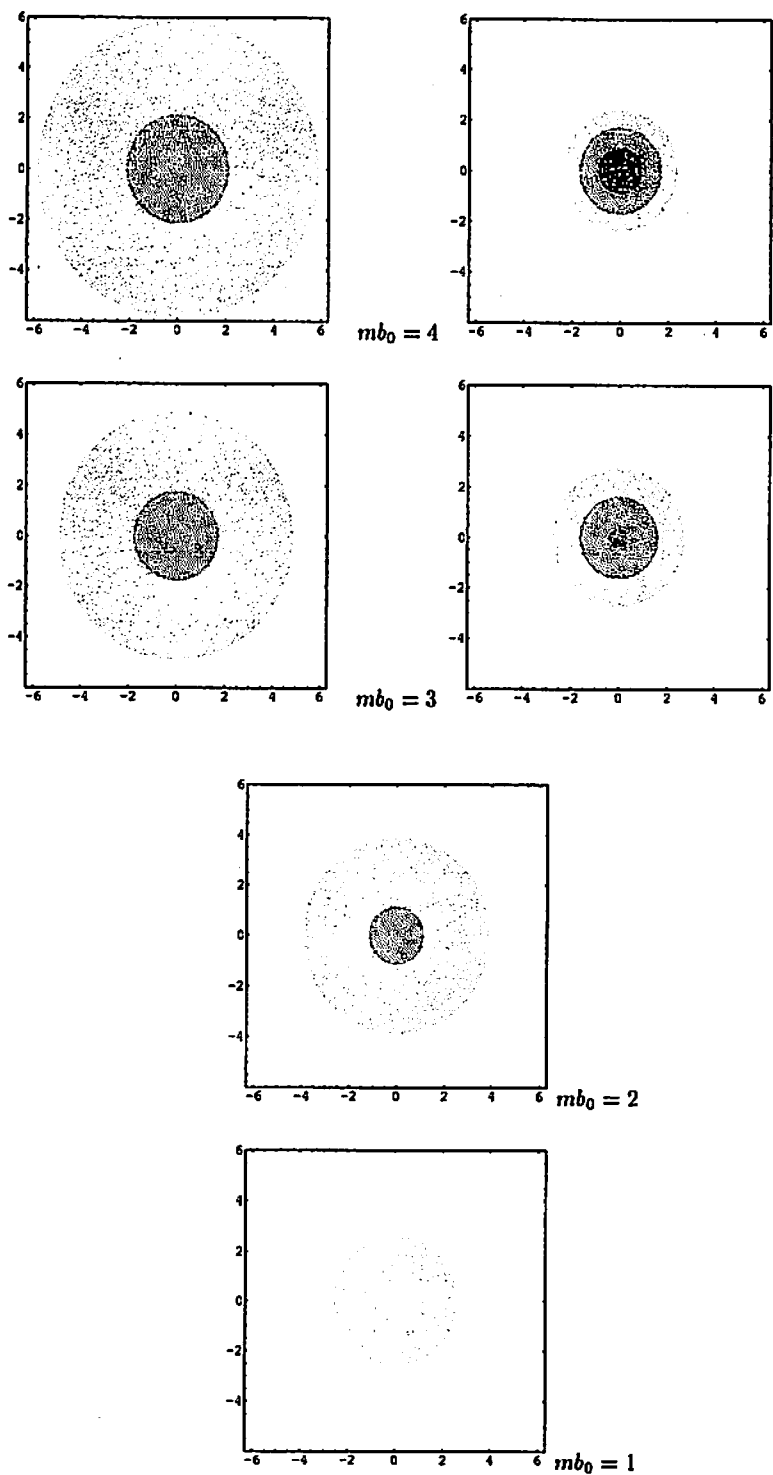


Fig 3: The evolutions of the fermion stars with the increase of the extra dimension. In the darkest circle, $n=2$ mode is most effective and $n=1$ mode in the middle one. The pale circle stands for no excited mode.

6 Conclusion

We have shown that considering K-K mode in the five dimensional theory, as the extra dimension becomes larger, the maximum mass more and more decreases. And also we have obtained two series of solutions. One is that the fermion star is getting larger and its central density lower with the increase of the extra dimension, that is, the larger and lighter star is formed. In these solutions, only $n=1$ mode is caused in the core of the stars. The other is that as the extra dimension grows, though the star once becomes larger and its central density lower for $mb_0=2$, it is getting smaller and smaller and its central density higher for the much larger dimension, namely, the smaller and heavier star is created. In these stars, $n=2,1$ modes is caused, the former is in the center of the stars and the latter is in the vast region including the core. As far as we examine, in this time, we have proved that the structure of the fermion stars depends on the scale of the fifth dimension, that is, the excited modes have effects to the inside of stars. Taking the scale of the fifth dimension much larger, however, we need to analyze the stability of stars explicitly. And, in this paper, we have imposed the periodic boundary condition on a wave function in the fifth dimension, while we can adopt the general one, that is: $\psi(\chi + 2\pi b) \sim e^{i\varphi}\psi(\chi)$. For the anti-periodic one, $\psi(\chi + 2\pi b) \sim -\psi(\chi)$, the effective mass of the fermion $M_{(h)}$ is $M_{(h)} = \sqrt{m^2 + (\frac{n\pm\frac{1}{2}}{b})^2}$, on which we are working, so, the details will be explained elsewhere. Meanwhile, we can apply our topics to cosmology in which we will suppose the time dependence of the fifth dimension and think over how the stars should be created in the time-dependent process. Although we took the zero temperature limit in this paper, we can treat the finite one too. Furthermore we will try to extend our argument into higher dimensional theories (the six dimensions, the ten dimensions and the others), and consider what the star made of the bulk matter in the brane world should be, in which we are going to research as the next theme.

References

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